

Dynamics of Robertson-Walker spacetimes with diffusion

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Abstract

We study the dynamics of spatially homogeneous and isotropic spacetimes containing a fluid undergoing microscopic velocity diffusion in a cosmological scalar field. After deriving a few exact solutions of the equations, we continue by analyzing the qualitative behavior of general solutions. To this purpose we recast the equations in the form of a two dimensional dynamical system and perform a global analysis of the flow. Among the admissible behaviors, we find solutions that are asymptotically de-Sitter both in the past and future time directions and which undergo accelerated expansion at all times.

1 Introduction

In a recent paper [4] one of us (SC) introduced a new type of fluid matter model in general relativity, in which the fluid particles undergo microscopic velocity diffusion in a cosmological scalar field. The energy-momentum tensor $T^{\mu\nu}$ and the current density J^μ of the fluid are given by the usual expressions

$$T^{\mu\nu} = \rho u^\mu u^\nu + p(g^{\mu\nu} + u^\mu u^\nu), \quad J^\mu = n u^\mu, \quad (1)$$

where ρ is the rest-frame energy density, p the pressure, u the four-velocity and n the particle density of the fluid. The diffusion behavior is imposed by postulating the equations

$$\nabla_\mu T^{\mu\nu} = \sigma J^\nu, \quad (2a)$$

$$\nabla_\mu (n u^\mu) = 0, \quad (2b)$$

where $\sigma > 0$ is the diffusion constant, which measures the energy gained by the particles per unit of time due to the action of the diffusion forces. Equation (2b) expresses the usual conservation of the total number of fluid particles, while (2a) is the diffusion equation. In fact it was shown in [3, 4] that (2a) is the (formal) macroscopic limit of a Fokker-Planck equation on the kinetic particle density, which is a standard model for diffusion dynamics taking place at the microscopic level on the particles velocity [15].

By projecting (2a) along the direction of u^μ and onto the hypersurface orthogonal to u^μ , we obtain the following equations on the matter fields:

$$\nabla_\mu (\rho u^\mu) + p \nabla_\mu u^\mu = \sigma n, \quad (3a)$$

$$(\rho + p) u^\mu \nabla_\mu u^\nu + u^\nu u^\mu \nabla_\mu p + \nabla^\nu p = 0, \quad (3b)$$

$$\nabla_\mu (n u^\mu) = 0. \quad (3c)$$

The system (3) on the matter variables may be completed by assigning an equation of state between the pressure and the energy density (barotropic fluid). In this paper we assume that

$$p = (\gamma - 1)\rho, \quad (4)$$

for some constant $2/3 < \gamma < 2$, so that in particular the fluid satisfies the strong and dominant energy conditions. The case $\gamma = 1$ gives rise to a dust fluid, while $\gamma = 4/3$ corresponds to a radiation fluid.

Since the energy-momentum is not divergence-free, the coupling of the Einstein equations with the diffusion equation (2a) is incompatible with the (contracted) Bianchi identity

$$\nabla^\mu \left(R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R \right) = 0. \quad (5)$$

This incompatibility can be resolved by postulating the existence of an additional matter field interacting with the fluid particles and restoring the local conservation of energy. This new matter field plays the role of the background medium in which the particles undergo diffusion.

The simplest model for this medium is a vacuum-energy source described by a cosmological scalar field, which leads to the following Einstein equations (in units $8\pi G = c = 1$):

$$R_{\mu\nu} - \frac{1}{2}g_{\mu\nu}R + \phi g_{\mu\nu} = T_{\mu\nu}. \quad (6)$$

The evolution equation on the scalar field ϕ determined by (6), the Bianchi identity (5), and the diffusion equation (2a) is

$$\nabla_\nu \phi = \sigma J_\nu. \quad (7)$$

Notice that when $\sigma = 0$, the model under study reduces to the Einstein-perfect fluid system with cosmological constant. The latter has important applications in cosmology, where, under appropriate symmetry assumptions, leads to the Λ CDM model, currently the most popular cosmological model of the universe. The analogous cosmological model with diffusion is the ϕ CDM model, in which Λ is replaced by the “variable cosmological constant” ϕ , representing a dark energy field interacting with dark matter by diffusion. This cosmological model has been studied in detail in [5], where in particular an upper bound on the diffusion constant σ has been estimated in order for the model to be compatible with the current cosmological observations. We emphasize that the interaction between the dark energy field and the matter fluid in the diffusion model is a natural and compelling consequence of the relativistic diffusion mechanism. Therefore the diffusion model considered here is fundamentally different from the other interacting models existing in the literature, where the interaction between dark energy and dark matter is supported only on a pure phenomenological basis; see [5] for a more detailed discussion on this important point. Further studies on the diffusion model described above can be found in [16, 17, 18].

In the present work we derive a complete characterization of the diffusion dynamics in a spatially homogeneous and isotropic spacetime. Using standard Robertson-Walker coordinates (t, r, θ, φ) , we write the metric as

$$ds^2 = -dt^2 + a(t)^2 \left[\frac{dr^2}{1 - kr^2} + r^2 d\Omega^2 \right], \quad d\Omega^2 = d\theta^2 + \sin^2 \theta d\varphi^2, \quad (8)$$

where $k = 0, \pm 1$ is the curvature parameter. The equations for the scale factor $a(t)$, the energy density $\rho(t)$ of the fluid and the cosmological scalar field $\phi(t)$ are

$$\dot{a} = Ha, \quad (9a)$$

$$\dot{\rho} = -3\gamma H\rho - \dot{\phi}, \quad (9b)$$

$$\dot{\phi} = -\sigma n_0 \left(\frac{a_0}{a(t)} \right)^3, \quad (9c)$$

$$\dot{H} = \frac{1}{3} \left[\phi - \left(\frac{3}{2}\gamma - 1 \right) \rho \right] - H^2, \quad (9d)$$

$$H^2 = \frac{1}{3}(\rho + \phi) - \frac{k}{a(t)^2}, \quad (9e)$$

where $a_0 > 0$, $n_0 > 0$ are the values of the scale factor and the particle density at the time $t = 0$. The behavior of solutions to the preceding system for $k = 0$ has been investigated numerically in [4]. In the present paper we provide analytical proofs for these numerical results and many others which were not discovered in [4]. For instance we show the existence of solutions to the Einstein equations (9) which exhibit accelerated expansion for all times, thereby providing a unified cosmological model for primordial inflation and late-time acceleration of the universe. Other interesting new types of solutions found in this paper include solutions for $k = -1$ which can be matched to anti-de-Sitter at a finite time in the past and solutions for $k = 1$ which describe an initially contracting, then expanding and finally recollapsing spacetime. These are only examples of the new and rich dynamics of the diffusion model discovered in the present paper.

The proof of our results is based on the application of techniques from the theory of finite dimensional dynamical systems, which is by now a standard tool in theoretical cosmology. An extensive discussion on this method for perfect fluid cosmological models without diffusion ($\sigma = 0$) and with $\Lambda = 0$ can be found in [20]. The dynamical systems approach has also been applied for the analysis of cosmological models with positive cosmological constant [11] and with matter sources other than fluids [6, 12]. When the cosmological constant is replaced by a minimally coupled scalar field, the case with exponential potential has been studied in [7]. Only recently an efficient dynamical systems approach has been introduced for monomial potentials [1, 2]. See also [8, 9, 19], and references therein, for other types of potential.

The rest of the paper is organized as follows. In Section 2 we solve the system (9) explicitly for the case where the scale factor $a(t)$ in the metric is linear on time. Moreover we point out the possible existence of solutions which can be matched to vacuum solutions in the past of some time t_0 and introduce a classification of solutions to (9) based on their asymptotic behavior toward the future and past time directions. The fact that solutions with such behaviors do exist is proved rigorously in Section 3. The proof is obtained by rewriting (9) using new dynamical variables, in such a way that the admissible self-similar asymptotic states of the system are represented by fixed points of the dynamical system in the new variables. The dynamical system in question is two dimensional but, as opposed to the diffusion-free case, it is not defined on a compact state-space. We remark that the new variables employed here are not obtained by normalizing the old ones with the Hubble function, as those used for instance in [16, 17, 18], and therefore we do not restrain our models to be forever expanding or contracting. This explains the considerably richer dynamics found in our investigation compared to the earlier available results on this model.

2 Classification of cosmological models

The initial data set for the system (9) consists of $(a_0, \rho_0, H_0, \phi_0)$, where a_0 , H_0 , ρ_0 and ϕ_0 satisfy (9e) at time $t = 0$, i.e.,

$$H_0^2 = \frac{1}{3}(\rho_0 + \phi_0) - \frac{k}{a_0^2}. \quad (10)$$

We remark that we do not necessarily identify $t = 0$ with the present time. For $k = 0, \pm 1$, we denote by $\mathcal{I}_k$ the three-dimensional manifold of admissible initial data, i.e.,

$$\mathcal{I}_k = \{(a_0, \rho_0, H_0, \phi_0) \in (0, \infty) \times (0, \infty) \times \mathbb{R} \times \mathbb{R} : (10) \text{ holds}\}.$$

Moreover we denote by q the deceleration parameter:

$$q = -\frac{a\ddot{a}}{\dot{a}^2} = -1 - \frac{\dot{H}}{H^2}. \quad (11)$$

We shall say that a solution of the Einstein equations (9) is expanding with acceleration if $H > 0$ and $q < 0$, and that it is contracting with acceleration if $H < 0$ and $q < 0$. Moreover, we say that a solution is singular at some time t_0 if $a(t_0) = 0$; we do not discuss here the question of whether the corresponding spacetime singularity is a curvature singularity or just a coordinate singularity.

The main purpose of this section is to introduce a classification of the solutions to the Einstein equations (9) based primarily on their asymptotic behavior toward the past and the future. We first introduce a number of exact solutions.

2.1 Exact solutions

Given the large variety of numerical solutions discovered in [4], it is unlikely that one can solve explicitly the system (9) for general initial data. However, in the particular case that the scale factor $a(t)$ is linear on time, we find the following explicit solution:

$$a(t) = a_0 + \delta_k t, \quad (12a)$$

$$\phi(t) = \frac{3\beta}{2\delta_k} a(t)^{-2}, \quad (12b)$$

$$\rho(t) = \frac{3\beta}{\delta_k(3\gamma - 2)} a(t)^{-2}, \quad (12c)$$

where we set

$$\beta = \frac{\sigma n_0 a_0^3}{3}$$

and δ_k is the real solution of the polynomial equation

$$\delta^3 + k\delta - \frac{3\beta\gamma}{2(3\gamma - 2)} = 0. \quad (13)$$

Note that $\delta_k > 0$, for all $k = 0, \pm 1$. In particular, for $k = 0$ the solution (12) becomes

$$a(t) = a_0 + \left(\frac{3\beta\gamma}{2(3\gamma - 2)} \right)^{1/3} t, \quad (14a)$$

$$\phi(t) = \left(\sqrt{\frac{3\gamma - 2}{\gamma}} \frac{3\beta}{2} \right)^{2/3} a(t)^{-2}, \quad (14b)$$

$$\rho(t) = \left(\frac{2}{\gamma} \right)^{1/3} \left(\frac{3\beta}{3\gamma - 2} \right) a(t)^{-2}. \quad (14c)$$

For $\sigma = 0$ the solution (14) reduces to the Minkowski spacetime, while the solution (12) for $k = -1$ reduces to the Milne spacetime. In contrast, there is no diffusion-free analogue of the solution (12) for $k = 1$.

All solutions defined by (12) are singularity free and forever expanding in the future, while in the past they become singular at the time $t_- = -a_0/\delta_k$. The future expansion takes place at a constant rate. In particular, the deceleration parameter defined by (11) is zero for the solutions (12). We also remark that the singularity at $t = t_-$ is a curvature singularity. In fact, computing the Ricci scalar of the solution (12) we obtain, using (6),

$$R = 4\phi + (4 - 3\gamma)\rho = \frac{9\beta\gamma}{\delta_k(3\gamma - 2)}a(t)^{-2} \rightarrow +\infty, \quad \text{as } a(t) \rightarrow 0^+.$$

2.2 Vacuum solutions

Next we discuss the vacuum solutions, i.e., solutions for which $\rho(t) = n(t) = 0$ for all times. Such solutions are important because they may act as future/past attractors of general solutions of the system (9). Since $\dot{\phi} = 0$ when $\rho = n = 0$, vacuum solutions correspond to maximally symmetric vacuum spacetimes with cosmological constant $\phi = \Lambda$. It is well known that any such spacetime is either de-Sitter (if $\Lambda > 0$), anti-de-Sitter (if $\Lambda < 0$), or Minkowski (if $\Lambda = 0$). We also recall that the coordinates (t, r, θ, φ) cover the full de-Sitter space for $k = 1$, while in all other cases they cover only a portion of (anti)-de-Sitter space. For more background on the geometry of de-Sitter/anti-de-Sitter we refer to [10, 13].

Letting $\rho = n = 0$ and $\phi = \Lambda$ in the system (9), we obtain that the scale factor of vacuum solutions satisfies

$$\frac{\ddot{a}}{a} - \left(\frac{\dot{a}}{a}\right)^2 = \frac{k}{a^2}, \quad a(0) = a_0, \quad \dot{a}(0) = H_0 a_0, \quad H_0 = \pm \sqrt{\frac{\Lambda}{3} - \frac{k}{a_0^2}}. \quad (15)$$

The explicit form of the solution to (15) depends on the values of the parameters k and Λ . Although the form of the scale factor in all various cases is well-known, we give it below for easy reference.

For $k = 0$ we must have $\Lambda \geq 0$ and the scale factor is given by

$$a(t) = a_0 e^{H_0 t}, \quad \text{where either } H_0 = \sqrt{\frac{\Lambda}{3}}, \quad \text{or } H_0 = -\sqrt{\frac{\Lambda}{3}}. \quad (16)$$

The corresponding spacetime is de-Sitter. For $\Lambda = 0$ (i.e., $H_0 = 0$) both solutions reduce to $a(t) = a_0$, and the corresponding spacetime is Minkowski.

For $k = 1$ we must have $\Lambda \geq 3/a_0^2$. The solutions are given by

$$a(t) = \sqrt{\frac{3}{\Lambda}} \cosh \left[\sqrt{\frac{\Lambda}{3}} t + \log \left[a_0 \left(\sqrt{\frac{\Lambda}{3}} + H_0 \right) \right] \right], \quad (17a)$$

$$\text{where either } H_0 = \sqrt{\frac{\Lambda}{3} - \frac{1}{a_0^2}}, \quad \text{or } H_0 = -\sqrt{\frac{\Lambda}{3} - \frac{1}{a_0^2}}. \quad (17b)$$

The two solutions coincide when $\Lambda = 3/a_0^2$ (i.e., $H_0 = 0$) and become

$$a(t) = a_0 \cosh\left(\sqrt{\frac{\Lambda}{3}} t\right). \quad (18)$$

The corresponding spacetime is de-Sitter.

For $k = -1$, the solutions are different according to the sign of Λ . For $\Lambda > 0$ we obtain

$$a(t) = \sqrt{\frac{3}{\Lambda}} \sinh\left[\sqrt{\frac{\Lambda}{3}} t + \log\left[a_0\left(\sqrt{\frac{\Lambda}{3}} + H_0\right)\right]\right], \quad (19a)$$

$$\text{where either } H_0 = \sqrt{\frac{\Lambda}{3} + \frac{1}{a_0^2}}, \quad \text{or} \quad H_0 = -\sqrt{\frac{\Lambda}{3} + \frac{1}{a_0^2}}. \quad (19b)$$

The corresponding spacetime is again de-Sitter.

For $k = -1$ and $\Lambda = 0$ we obtain the Milne solutions:

$$a(t) = a_0(1 + H_0 t), \quad \text{where either } H_0 = \frac{1}{a_0}, \quad \text{or} \quad H_0 = -\frac{1}{a_0}. \quad (20)$$

For $-3/a_0^2 \leq \Lambda < 0$, the solutions to (15) are given by

$$a(t) = \sqrt{\frac{3}{|\Lambda|}} \sin\left[\operatorname{arccot}\left(\sqrt{\frac{3}{|\Lambda|}} H_0\right) + \sqrt{\frac{|\Lambda|}{3}} t\right], \quad (21a)$$

$$\text{where } H_0 = \pm \sqrt{\frac{\Lambda}{3} + \frac{1}{a_0^2}}, \quad (21b)$$

and for $H_0 = 0$ they both reduce to

$$a(t) = \sqrt{\frac{3}{|\Lambda|}} \cos\left(\sqrt{\frac{|\Lambda|}{3}} t\right). \quad (22)$$

The corresponding spacetime is anti-de-Sitter.

2.3 Vacuum matching solutions

A remarkable difference with the diffusion-free case is the possibility that the energy density ρ vanished at some finite time t_0 while the scale factor a is still regular. This possibility arises because $\rho = 0$ is not a solution of (9b). By a time translation we may set $t_0 = 0$ for any given such type of solution and therefore we assume that

$$\rho(0) = 0, \quad a(0) > 0, \quad |H(0)| < \infty.$$

By (9b), $\dot{\rho}(0) > 0$, hence $\rho(0) < 0$ for $t < 0$. To avoid this unphysical region of spacetime, we replace it with either the de-Sitter or the anti-de-Sitter, depending on the sign of $\phi(0)$. This extended spacetime is singularity free in the past and in the region $t \leq 0$ covered by the coordinates (t, r, θ, ϕ) , the scale factor is given by one of the expressions (16), (17), (19), depending on the value of k , $\phi(0) = \Lambda$ and $H_0 = H(0)$. The matching at $t = 0$ is C^2 in the scale factor $a(t)$ and C^0 in the scalar field $\phi(t)$ and in the energy density $\rho(t)$.

2.4 Classification of general solutions

To conclude this section we introduce a classification of general solutions to (9) based on their asymptotic behavior toward the past and future time directions. In particular, we say that a solution is of type **A** if it becomes singular at some finite time in the past, while in the future it is singularity free and asymptotically de-Sitter. A solution is said to be of type **B** if it can be matched to a de-Sitter solution at some finite time in the past, while in the future it is singularity free and asymptotically de-Sitter. Solutions of type **C** are those which can be matched to a vacuum solution at some finite time in the past and which become singular at some finite time in the future. Finally, a solution of type **D** is a solution that becomes singular at finite time in both time directions.

The main result of this paper can be formally stated as follows: For all $k = 0, \pm 1$, there exists four disjoint three-dimensional submanifolds of initial data $\mathcal{A}_k \subset \mathcal{I}_k$, $\mathcal{B}_k \subset \mathcal{I}_k$, $\mathcal{C}_k \subset \mathcal{I}_k$, $\mathcal{D}_k \subset \mathcal{I}_k$ such that if the initial data belong to $\mathcal{A}_k$, the corresponding solution of (9) is of type **A**, if the initial data belong to $\mathcal{B}_k$, the corresponding solution is of type **B**, etc. Moreover solutions which are not launched by initial data in the set $\mathcal{A}_k \cup \mathcal{B}_k \cup \mathcal{C}_k \cup \mathcal{D}_k$ are atypical, that is to say, they correspond to initial data forming a two-dimensional submanifold of $\mathcal{I}_k$.

3 Qualitative dynamics of solutions

For an introduction to dynamical systems theory see [14]. Let

$$D = \sqrt{H^2 + \frac{1}{a^2}} \quad (23)$$

and define a new time variable τ by

$$\frac{d}{d\tau}(\cdot) = \frac{1}{D} \frac{d}{dt}(\cdot);$$

in the following we use the notation $(\cdot)' = \frac{d}{d\tau}(\cdot)$. We also introduce the dimensionless variables

$$H_D = \frac{H}{D}, \quad M_D = \frac{1}{aD}, \quad \Omega_D = \frac{\rho}{3D^2}, \quad Y_D = \frac{\phi}{3D^2}, \quad X_D = \frac{\dot{\phi}}{3D^3}. \quad (24)$$

These variables satisfy the algebraic constraints

$$H_D^2 + M_D^2 = 1, \quad X_D + \beta M_D^3 = 0, \quad H_D^2 = \Omega_D + Y_D - k M_D^2, \quad (25)$$

where we recall that $\beta = \sigma n_0 a_0^3/3$. The first equation in (25) follows by the definition of the new variables, while the second and third equations correspond respectively to (9c) and (9e).

The evolution equation for the dimensional variable D is given by

$$D' = -H_D D(1 + qH_D^2), \quad (26)$$

while the remaining variables satisfy

$$\begin{aligned} H_D' &= -qH_D^2(1 - H_D^2), \\ M_D' &= qH_D^3 M_D, \\ \Omega_D' &= -X_D + 2H_D \Omega_D(1 + qH_D^2 - \tfrac{3}{2}\gamma), \\ Y_D' &= X_D + 2H_D Y_D(1 + qH_D^2), \\ X_D' &= 3qH_D^3 X_D, \end{aligned}$$

where

$$qH_D^2 = -Y_D + \left(\frac{3}{2}\gamma - 1\right)\Omega_D. \quad (27)$$

Note that the equation on D decouples from the rest of the system; this is due to the fact that D is dimensional, while the other variables have no physical dimension (in our units).

From the constraint equations (25), the dimension of the dynamical system can be reduced by three, obtaining a 2-dimensional dynamical system. We choose to work with the variables (Y_D, H_D) . Once these are known, the remaining variables can be determined globally from the constraints. The reduced dynamical system we shall be working with is then given by

$$H_D' = [Y_D + (1 - \tfrac{3}{2}\gamma)\Omega_D](1 - H_D^2), \quad (28a)$$

$$Y_D' = -\beta(1 - H_D^2)^{\frac{3}{2}} + 2H_D Y_D [1 - Y_D - (1 - \tfrac{3}{2}\gamma)\Omega_D], \quad (28b)$$

where

$$\Omega_D = H_D^2 + k(1 - H_D^2) - Y_D. \quad (28c)$$

As opposed to the standard diffusion-free case (which is obtained by setting $\beta = 0$ in the equations above), the variable Ω_D is not bounded and $\Omega_D = 0$ is not an invariant boundary. However, since

$$(\Omega_D')|_{\Omega_D=0} = -X_D = \beta(1 - H_D^2)^{3/2} > 0,$$

the curve $\Omega_D = 0$ acts as “semipermeable membrane”: the flow can cross this line only in one direction. In particular, the region $\Omega_D > 0$ is future invariant and if an orbit in the region $\Omega_D > 0$ intersects the vacuum line $\Omega_D = 0$ in the past, then Ω_D is negative for all earlier times along this orbit. As discussed in Section 2.3, the solutions of the original system (9) corresponding to these orbits can be matched to a suitable vacuum solution at the time when the boundary Ω_D is crossed. We emphasize once again that this matching is not possible in the diffusion-free case, because $\Omega_D = 0$ is an invariant boundary when $\beta = 0$.

By the preceding remarks, we only need to worry about the region $\Omega_D > 0$, which in terms of the variables (Y_D, H_D) means

$$\Omega_D > 0 \Leftrightarrow Y_D < H_D^2 + k(1 - H_D^2). \quad (29)$$

The state-space $\mathcal{X}$ for the reduced dynamical system (28) is then

$$\mathcal{X} = \{(Y_D, H_D) \in \mathbb{R} \times (-1, 1) : \Omega_D > 0\}.$$

Moreover the dynamical system admits a continuous extension on the boundary $H_D = \pm 1$.

Our next goal is to study the qualitative behavior of the flow of the dynamical system (28) and to present the physical interpretation of this analysis in terms of solutions of the Einstein equations (9) and their asymptotic behavior.

3.1 Fixed points

The dynamical system (28) possesses five fixed points, four of which are located on the boundary ($H_D = \pm 1$) and one in the interior, see Table 3.1. The interior fixed point $S_{(k)}$ is associated to the self-similar solutions (14) which have been characterized in Section 2, with $a(t)$ being a linear function on time. Since the other fixed points are located at the boundary of the state-space, they no longer correspond to exact solutions of (9), but to limiting states when one or more variables take an extreme value.

The main properties of the fixed points of the dynamical system (28) are the following.

$S_{(k)}$: At this fixed point we have

$$Y_D = \left(1 - \frac{2}{3\gamma}\right) \left(k + (1 - k) \frac{\delta_k^2}{1 + \delta_k^2}\right)$$

and

$$H_D = \frac{\delta_k}{\sqrt{1 + \delta_k^2}},$$

where δ_k is the (real, positive) solution of (13) depending on k . Hence this fixed point identifies the self-similar solutions (12). The fixed point $S_{(k)}$ is hyperbolic and a simple local stability analysis shows that $S_{(k)}$ is a saddle fixed point.

$F_{\pm}$: At these fixed points we have $Y_D = 0$ and $H_D = \pm 1$. The fixed point F_+ is a source of interior orbits. To see this, consider the flow induced by (28) on the boundary $H_D = 1$. This flow is described by the one-dimensional dynamical system

$$Y'_D = 3\gamma Y_D(1 - Y_D).$$

As $Y'_D > 0$ for $Y_D \in (0, 1)$ and $Y'_D < 0$ for $Y_D < 0$, the fixed point F_+ repels the orbits on the boundary $H_D = 1$. Moreover, using that

$$\left(\frac{d}{d\tau} \log(1 - H_D^2)\right)_{(Y_D, H_D)=(0,1)} = 3\gamma - 2 > 0,$$

we obtain that the variable H_D is decreasing along interior orbits approaching F_+ . It follows that the fixed point F_+ repels interior orbits as well and thus that it is a source of interior orbits. We remark that since F_+ lies on the boundary of the state space, this result cannot be obtained by linearizing the system around this fixed point. A similar argument shows that F_- is a sink of interior orbits. From the constraints (25) we obtain $M_D = X_D = 0$ and $\Omega_D = 1$. Substituting in (26) and going back to the original time variable t we obtain

$$\dot{D} = \mp \frac{3}{2} \gamma D^2. \quad (30)$$

It follows that along the solutions of the system (9) which correspond to the interior orbits converging to (resp. emanating from) the fixed point F_- (resp. F_+), the variable $D(t)$ blows-up in finite time toward the future (resp. past). For any such solution we may choose the origin of time so that the singularity occurs at $t = 0$ and integrating (30) we find

$$D(t) \sim \pm \frac{\frac{2}{3\gamma}}{t}, \quad \text{as } t \rightarrow 0^\pm.$$

Since $H_D \rightarrow \pm 1$ implies $H(t) \sim \pm D(t)$, we obtain

$$H(t) \sim \frac{2}{t}, \quad \text{as } t \rightarrow 0^\pm. \quad (31)$$

Thus orbits that emanate from the fixed point F_+ (resp. converge to F_-) identify solutions of the Einstein equations which behave like the expanding (resp. contracting) diffusion-free perfect fluid solution with zero cosmological constant (i.e., the Friedmann-Lemaître solution) in the limit toward the past (resp. future) singularity.

dS_±: At these fixed points we have $Y_D = 1$ and $H_D = \pm 1$, hence (25) gives $M_D = X_D = \Omega_D = 0$. An argument similar to the one used for the fixed point F_+ shows that the fixed point dS_- is a source, while dS_+ is an attractor of interior orbits. By (26) we obtain $D(t) = c$, which implies that $H(t) = \pm c$, where c is a positive constant. It follows that the orbits that converge to the fixed point dS_+ (resp. emanate from dS_-) correspond to solutions of the Einstein equations which are singularity free in the future (resp. past) and behave like the expanding (resp. contracting) de-Sitter solution for late (resp. early) times for $t \rightarrow +\infty$ (resp. $t \rightarrow -\infty$).

3.2 Analysis of the flow

We begin by pointing out a few general properties of the flow of the dynamical system (28), which are illustrated in Figure 1. As already mentioned, the region $\Omega_D > 0$ is future invariant. Similarly, since

$$(Y_D')|_{Y_D=0} = -\beta(1 - H_D^2)^{3/2} < 0,$$

Fixed point	Y_D	H_D
dS ₋	1	-1
dS ₊	1	1
F ₋	0	-1
F ₊	0	1
S _k	$(1 - \frac{2}{3\gamma})(k + (1 - k)\frac{\delta_k^2}{1+\delta_k^2})$	$\frac{\delta_k}{\sqrt{1+\delta_k^2}}$

Table 3.1: Fixed points of the dynamical system (28) in the state-space $\overline{\mathcal{X}}$. δ_k is the solution of (13).

the region $Y_D < 0$ is future invariant. Hence if an orbit crosses the line $Y_D = 0$ from the right to the left, it will remain on the left region for all future times. Furthermore, in the region $Y_D < 0$ equation (28a) implies that H_D is monotone decreasing and that the subregion $H_D < 0$ is future invariant. It follows easily that the ω -limit of all orbits that cross the line $Y_D = 0$ from the right to the left is the fixed point F₋. Moreover if an orbit is entirely contained in the region $Y_D < 0$, then the α -limit of this orbit is the fixed point F₊.

Since the fixed point S_(k) is a saddle, it is the ω -limit set of two interior orbits and the α -limit set of two other interior orbits. These orbits divide the state-space in four regions, which we denote by **A**, **B**, **C**, and **D**; see Figure 1. Each orbit in the region **A** corresponds to solutions of the Einstein equations of type **A**, as defined in Section 2.4 and similarly for the other regions. The shadowed region in the pictures in Figure 1 identifies the region of accelerated expansion (if $H_D > 0$) and contraction (if $H_D < 0$), i.e., the region where $q < 0$, which is

$$\left(1 - \frac{2}{3\gamma}\right)(H_D^2 + k(1 - H_D^2)) < Y_D < (H_D^2 + k(1 - H_D^2)),$$

cf. (27) and (29). We continue by studying the qualitative behavior of the orbits in these regions separately for the cases $k = 0$, $k = -1$ and $k = 1$.

Case $k = 0$

The qualitative behavior of orbits in the various regions of the state-space in the case $k = 0$ is depicted in Figure 2 and can be described as follows.

The α - and ω -limit set of the orbits in the region **A** are the fixed points F₊ and dS₊, respectively. Hence the solutions of the system (9) corresponding to these orbits have a singularity in the past, toward which they behave like the expanding Friedman solution, while in the future they are singularity free and asymptotically de-Sitter. Moreover $H_D > 0$ in the region **A** and all orbits enter the shadowed region before converging to dS₊. Therefore the corresponding solutions of the Einstein equations are forever expanding and, after some finite time, the expansion becomes accelerated.

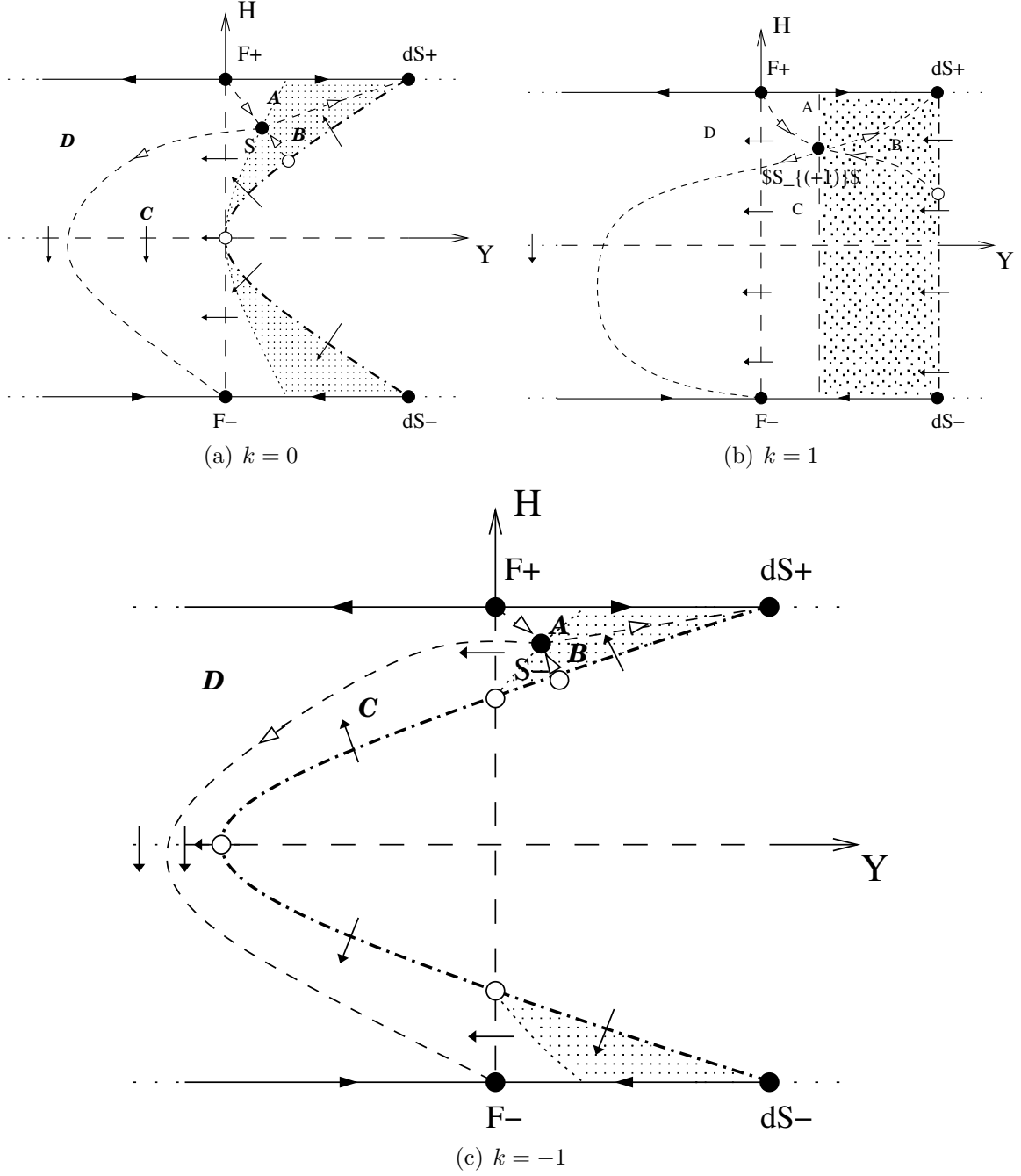


Figure 1: Partition of the state-space

All orbits in the region **B** converge to the fixed point dS_+ , hence the corresponding solutions of the Einstein equations are singularity free and asymptotically de-Sitter in the future time direction. In the past direction, orbits are forced to intersect the vacuum line $\Omega_D = 0$. In terms of solutions to the system (9) this means that, letting $t = 0$ be the time of

intersection, the scale factor can be continued to (16) (with $H_0 > 0$) for $t \leq 0$. The resulting cosmological model is then given by the de-Sitter spacetime for $t \leq 0$. Furthermore, since $H_D > 0$ and the region **B** is completely shadowed, solutions corresponding to orbits in this region undergo accelerated expansion *for all times*.

Orbits in the region **C** converge to the fixed point F_- in the future, which means that the corresponding solutions of the Einstein equations possess a singularity in the future toward which they behave like the contracting Friedman solution. Toward the past, orbits in the region **C** intersect the vacuum line $\Omega_D = 0$. It is worth to distinguish the behavior of the corresponding solutions of the system (9) into three different types. First we point out the existence of two special orbits in the region **C**, which we call c_1 and c_2 . The orbit c_1 passes through the point $(Y_D, H_D) = (0, 0)$ (which corresponds to Minkowski), while the orbit c_2 originates from dS_+ in the direction tangent to $\Omega_D = 0$ at that point. The orbits c_1, c_2 divide the region **C** into three invariant regions, see Figure 2(c). Solutions corresponding to orbits below c_2 are singularity free and asymptotically de-Sitter in the past. Moreover they are forever contracting, because $H_D < 0$. The orbits above c_2 intersect the vacuum line in the interior of the state-space (i.e., at finite time), and thus the corresponding solutions of the Einstein equations can be continued by de-Sitter prior to this time (i.e., by (16) with $H_0 < 0$). The difference between an orbit on the left of c_1 and an orbit on the right of c_1 is that the former corresponds to a solution which is initially expanding and then recollapsing (because the sign of H_D changes from positive to negative along the orbit), while the latter corresponds to a forever contracting solution (since $H_D < 0$ along the orbit.)

Finally, the α -limit set of orbits in the region **D** is the point F_+ , while the ω -limit set is F_- . This means that the corresponding solutions of the system (9) become singular at finite time in both time directions. Moreover, since the sign of H_D changes from positive to negative along all orbits, the solutions of the Einstein equations associated to the orbits in the region **D** are initially expanding, until they reach a stage of maximum extension, and then they recollapse into the future singularity.

Case $k = -1$

The qualitative behavior of orbits for $k = -1$ is very similar to the case $k = 0$, as it can be seen in Figure 3. We merely point out that in the region **C**, depicted in Figure 3(c), the matching of the scale factor at $\Omega_D = 0$ is either with de-Sitter (i.e., (19)), or with Milne (i.e., (20)), or with anti-de-Sitter (i.e., (21)). In particular, there exists two special orbits, which we call respectively c_3 and c_4 , intersecting the vacuum line in the expanding and contracting Milne solution. Both these orbits converge to F_- in the future and divide the region **C** into three parts. The solutions of the Einstein equation corresponding to orbits in the interior part delimited by c_3 and c_4 are matched to the anti-de-Sitter solution, while the solutions corresponding to the two exterior regions above c_2 are matched with de-Sitter. Below c_2 , the solutions are asymptotic to de-Sitter in the past.

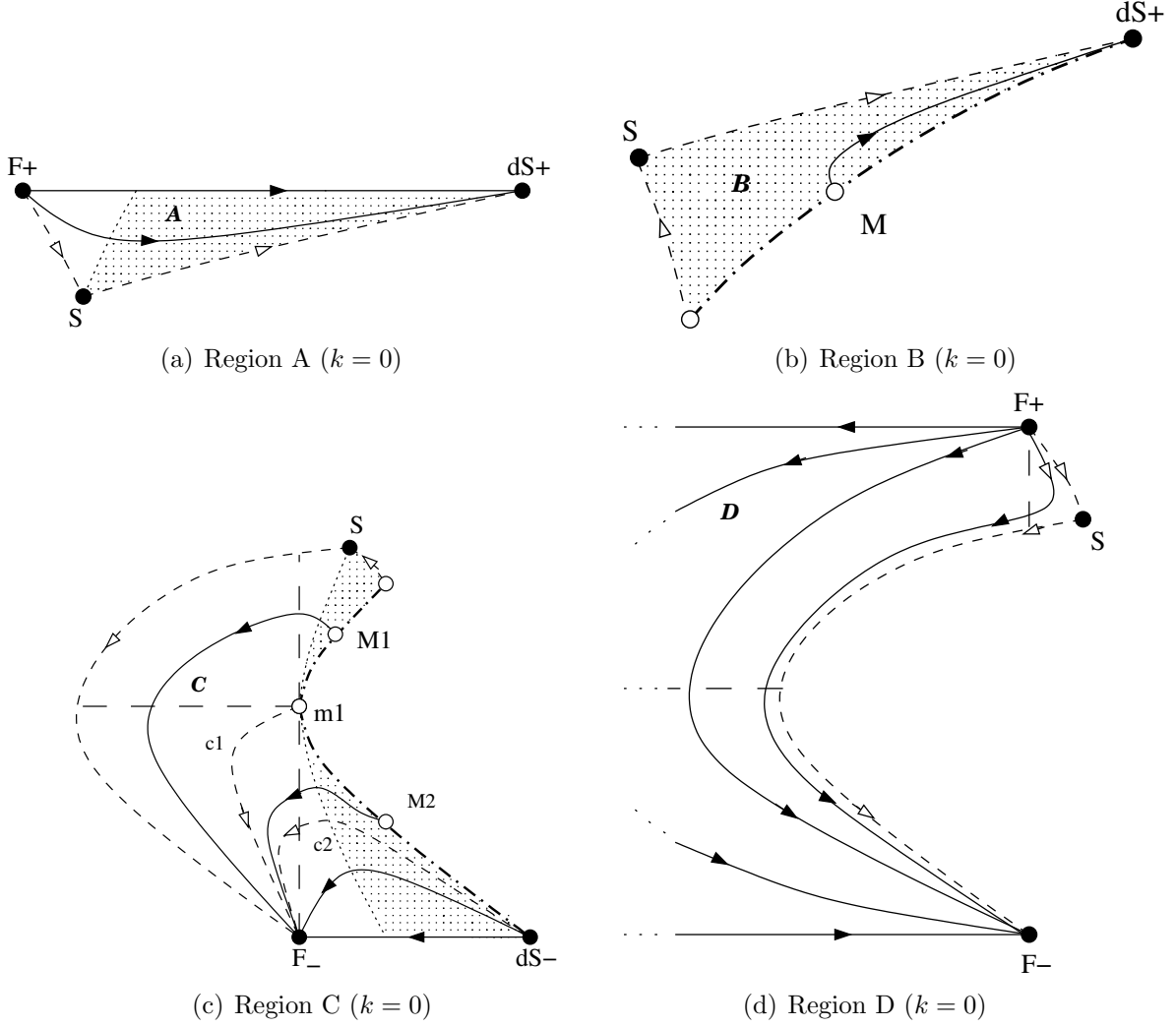


Figure 2: Phase portrait in the various regions of state-space ($k = 0$)

Case $k = 1$

The qualitative behavior of the orbits in regions **A**, **B** and **D** for this case is similar to the previous cases $k = 0$ and $k = -1$, as shown in Figure 4. As for region **C**, depicted in Figure 4(c), the matching at $\Omega_D = 0$ of the orbits above the orbit c_2 is with de-Sitter (i.e., (19)), while below c_2 the solutions are asymptotic to de-Sitter in the past. In this case, there are two special orbits, c_5 and c_6 , intersecting the vacuum line (on the boundary of the phase space) in the expanding and contracting de-Sitter solutions, respectively. Both these orbits converge to F_- in the future and divide the region **C** into three parts. The interior part delimited by c_5 and c_6 consists of orbits that correspond to solutions which are initially contracting, then expanding and finally recollapsing (because the sign of H_D changes from negative to positive and then from positive to negative along the orbits). The orbits between

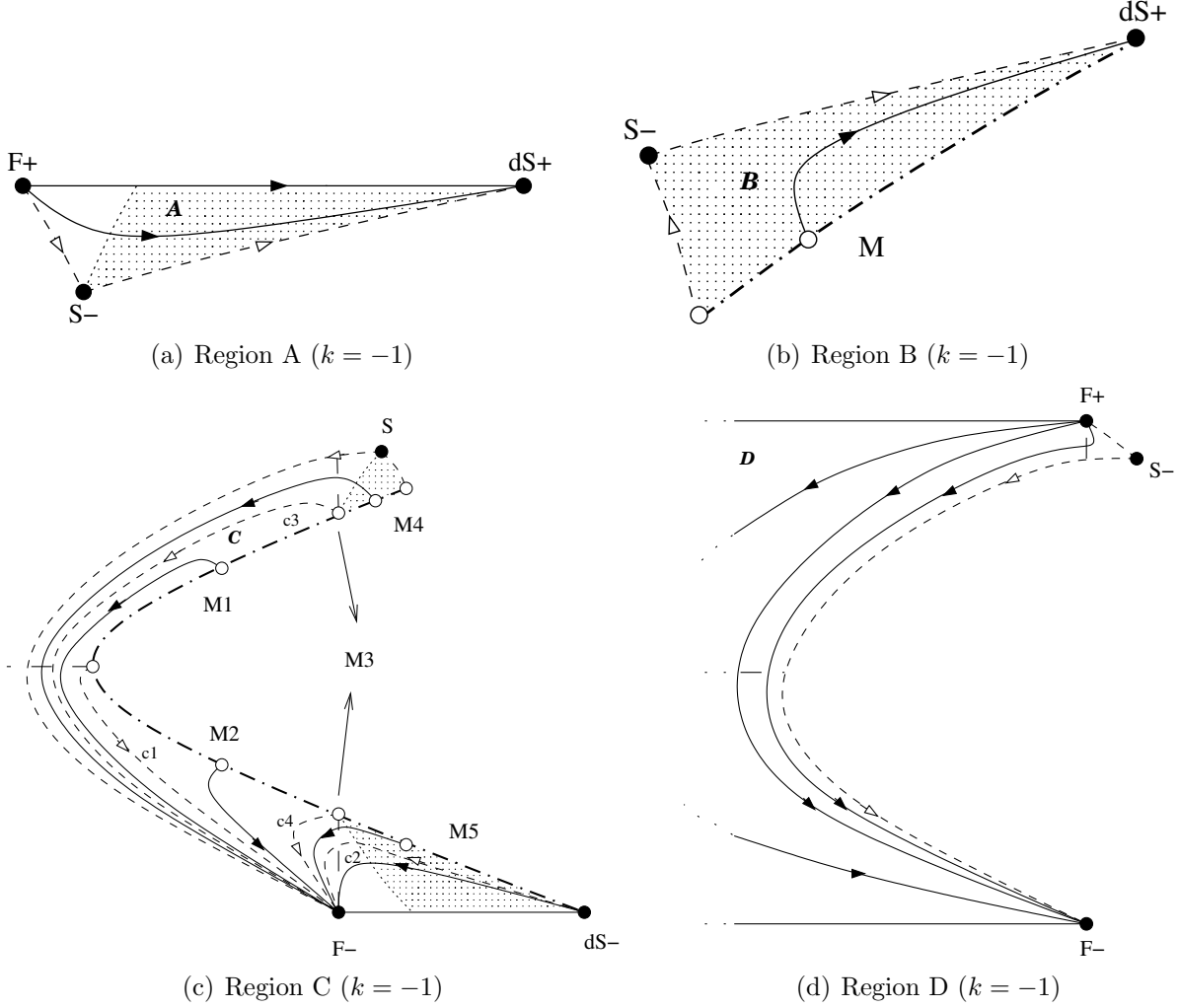


Figure 3: Phase portrait in the various regions of state-space ($k = -1$)

c_2 and c_6 are forever contracting (since $H_D < 0$ along the orbits). The orbits above c_5 are initially expanding and then collapsing (because the sign of H_D changes from positive to negative along the orbits).

4 Summary

In this work we considered a cosmological model based on the Einstein equations and the assumption that the matter content of the universe is described by a fluid undergoing diffusion in a cosmological scalar field. The scalar field can be identified with the dark energy field. We take spacetime to have a Robertson-Walker line element, so that the model studied here is spatially homogeneous and isotropic. The matter field variables are solutions of a non

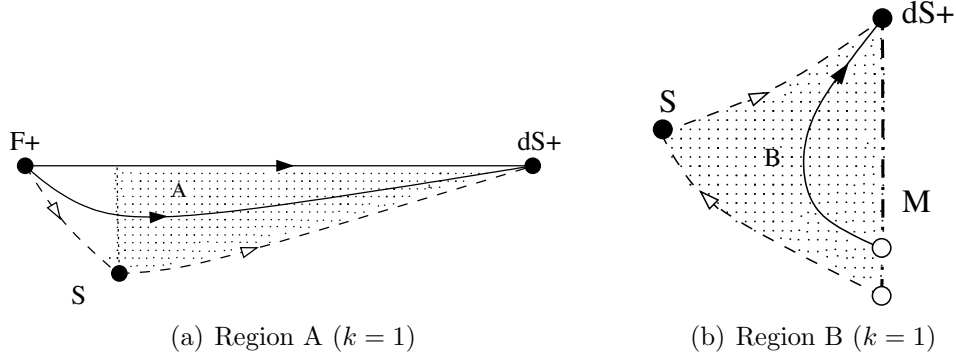
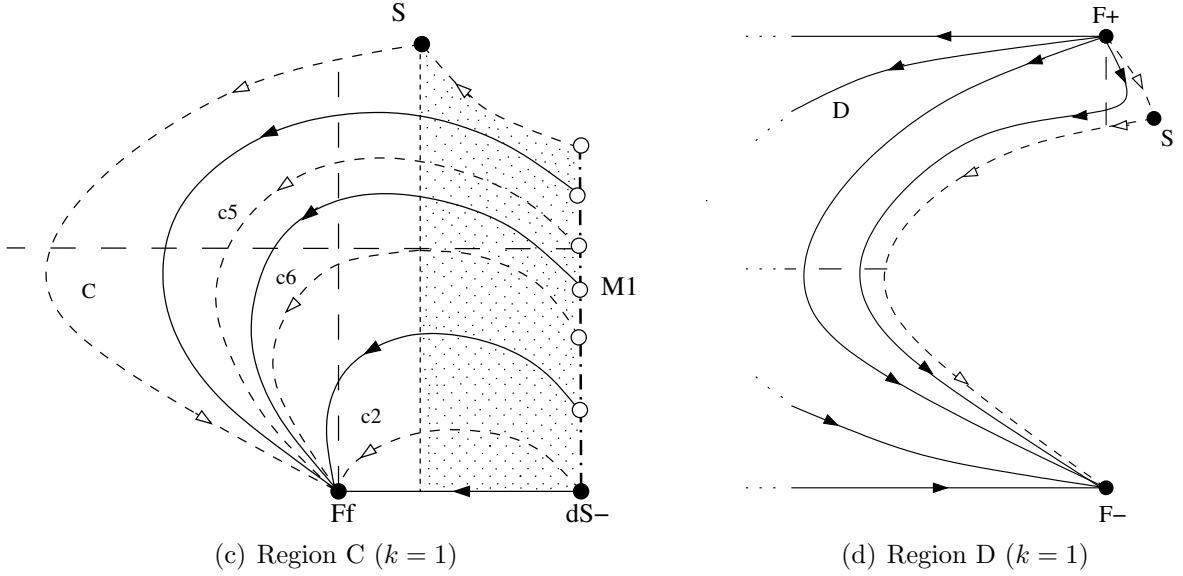


Figure 4: Phase portrait in the various regions of state-space ($k = 1$)

linear system of ordinary differential equations. We were able to obtain all solutions in which the scale factor is linear on time. In order to understand the dynamical properties of general solutions, we rewrite the system in terms of normalized (dimensionless) dynamical variables. We have shown that typical solutions of the Einstein equations can be classified according to their past and future asymptotic behavior into four classes, which we called **A**, **B**, **C**, **D**. Solutions of type **A** describe a forever expanding universe having a Big Bang singularity in the past and which is asymptotic to de-Sitter toward the future; type **B** solutions describe a singularity-free cosmological model which at some finite time in the past can be matched to de-Sitter. Solutions of this type are forever expanding with acceleration, and since the scalar field is monotonically decreasing, the future asymptotic de-Sitter state has a smaller value of the cosmological constant than the past de-Sitter matching solution. Solutions of type **C** and **D** have a Big Crunch singularity in the future, due to the scalar field becoming negative and thus mimicking a negative cosmological constant, but while solutions of type **C** have a Big Bang singularity in the past, those of type **D** are singularity free and can be matched

at finite time (or approach asymptotically) either de-Sitter or Anti-de-Sitter depending on the value of the curvature parameter k .

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